Supernova Bound on keV-mass Sterile Neutrinos

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Based on *G.G. Raffelt & S.Z., Phys.Rev.D 83, 093014 (2011)*

Workshop CIAS Meudon 2011

Outline

- Introduction
- Sterile neutrinos in SN cores
- Energy-loss rates and SN bounds
- Sterile neutrinos and SN explosions
- Summary

Introduction

What are sterile neutrinos?

- ✓ Standard Model (SM) Singlet
- ✓ Only mixing with SM *ν*'s

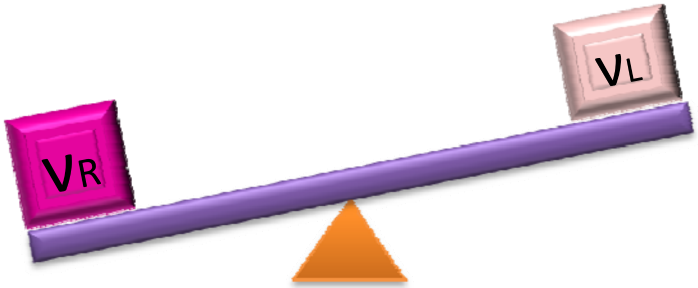
How to describe them?

- ➢ Mass m_s (eV, keV, MeV, …)
- ➢ Mixing angle ϑ with SM *ν*'s

A Typical Example: **Right-handed neutrinos in seesaw models**

Minkowski, 77'; Yanagida, 79'; Gell-mann, Ramond, Slansky, 79'

$$m_\nu \approx m_{\rm D} M_{\rm R}^{-1} m_{\rm D}^{\rm T}$$



Other Examples:

Mirror neutrinos, Berezhani & Mohapatra, 95'; Goldstinos, Chun et al., 96'; Modulinos, Benakli & Smirnov, 97'; ……

Introduction: *why keV-mass sterile neutrinos?*

Mass

Motivated by observations and experimental anomalies

GeV — Matter-antimatter asymmetry, ν masses, …

Asaka, Shaposhnikov, 05'; …

MeV — To reconcile LSND, MiniBooNE, and KARMEN

Gninenko, 11'

keV — Warm Dark Matter, Pulsar kicks, SN Explosions

Dodelson, Widrow, 94'; Shi, Fuller, 99'; Dolgov, Hansen, 02'; Fuller, Kusenko, Mocioiu, Pascoli, 03'; Hidaka, Fuller, 06' ; …

eV — LSND, MiniBooNE, Reactor ν anomaly, SBL,…

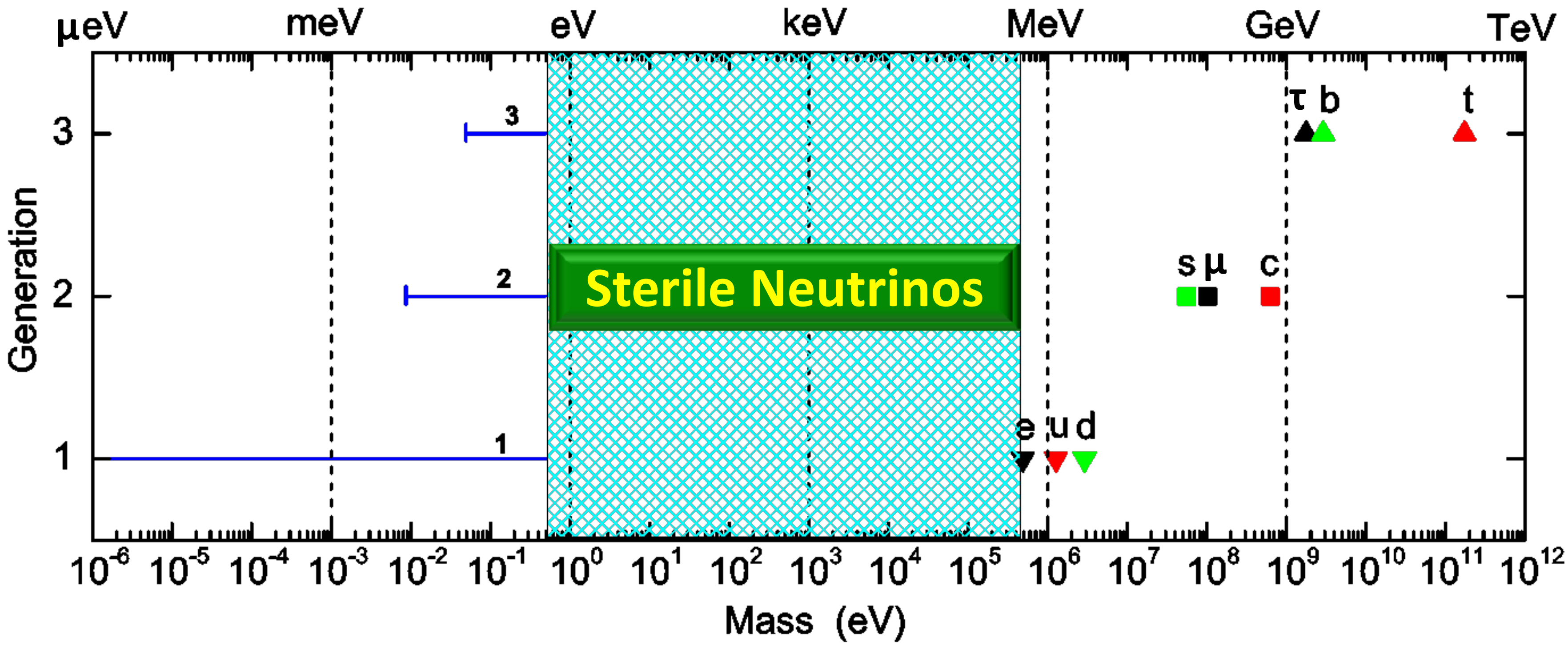
meV — 'Solar ν anomaly': non-observation of upturn

de Holanda, Smirnov, 04', 10'

Introduction: *why keV-mass sterile neutrinos?*

Xing, 11'

Motivated by the 'desert' in the fermion mass spectrum



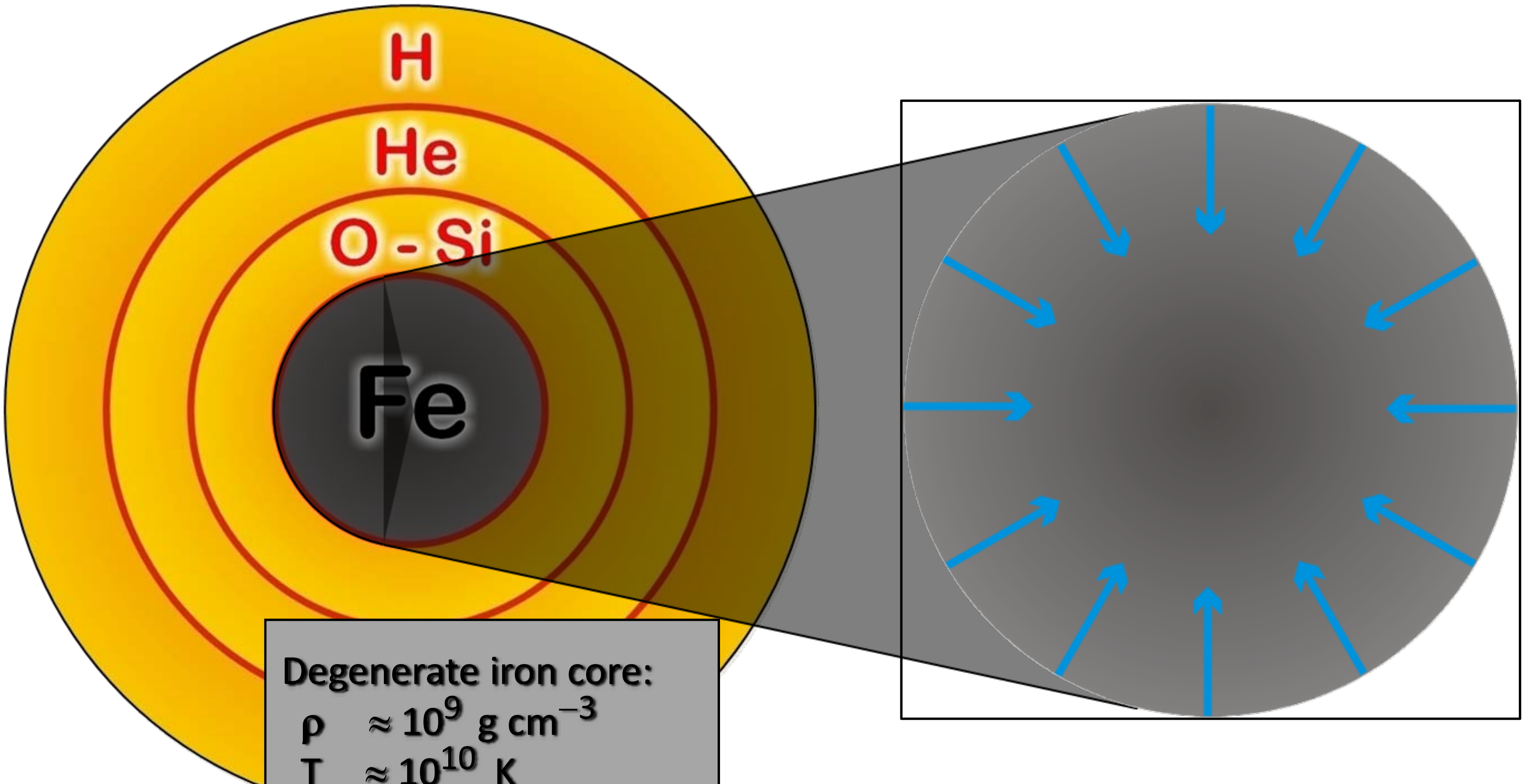
Constraints on keV-mass sterile neutrinos: m_s & ϑ

Relic abundance, Structure formation, X-ray observations, ……

What is the SN bound on keV-mass sterile neutrinos?

Introduction: *Supernova Explosions*

Onion structure | **Collapse (implosion)**



Degenerate iron core:

$\rho \approx 10^{9}\ \mathrm{g\ cm^{-3}}$

$T \approx 10^{10}\ \mathrm{K}$

$M_{Fe} \approx 1.5\ M_{sun}$

$R_{Fe} \approx 8000\ \mathrm{km}$

Introduction: *Supernova Explosions*

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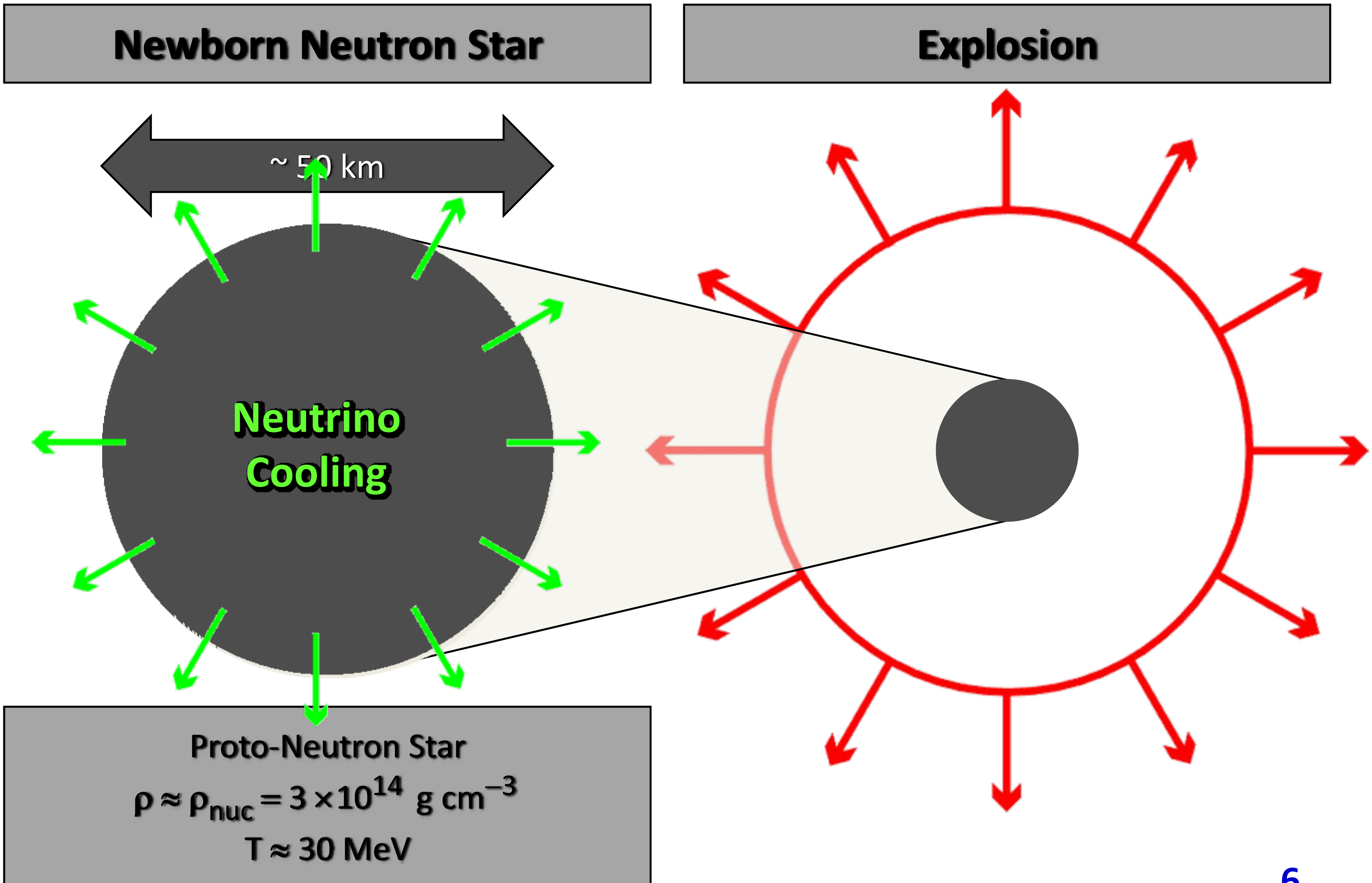
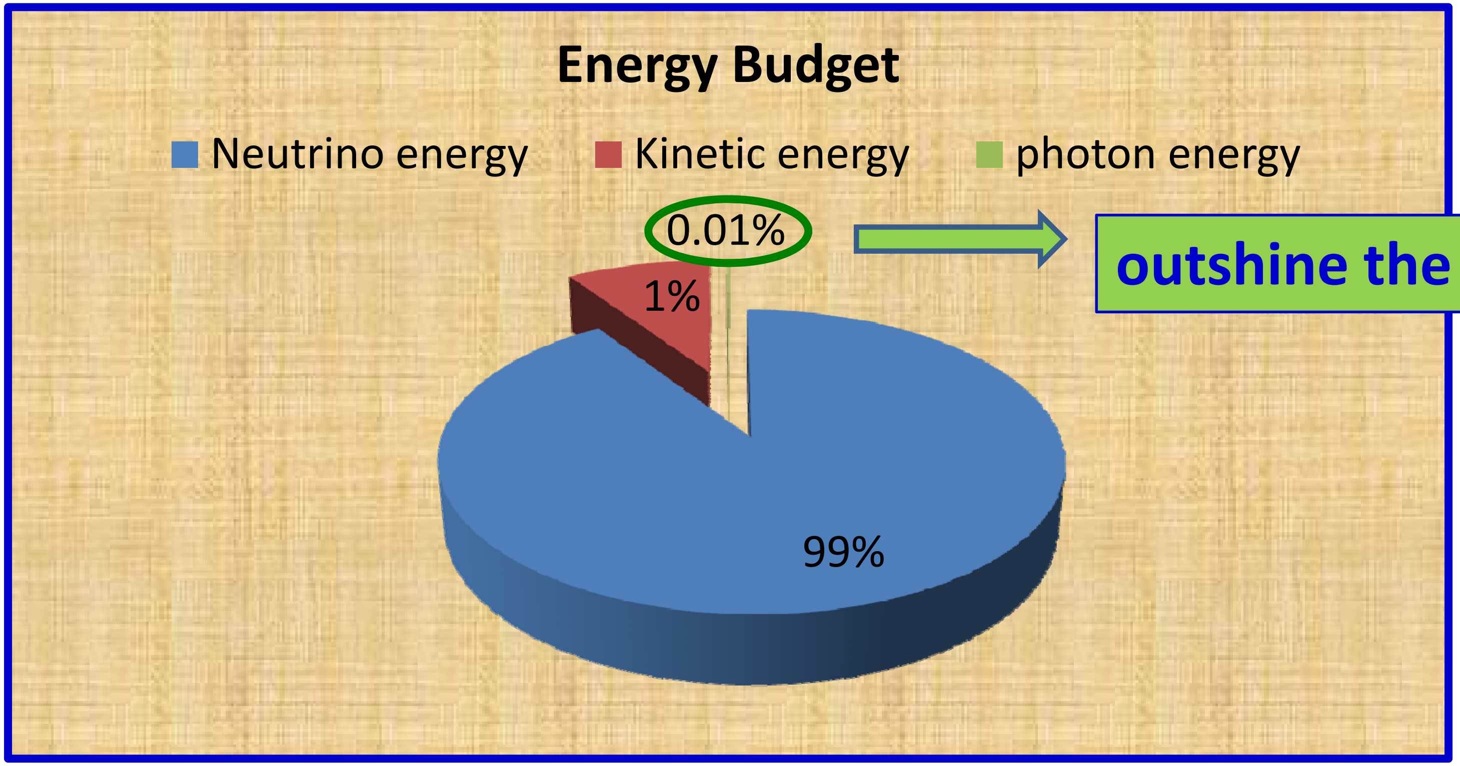


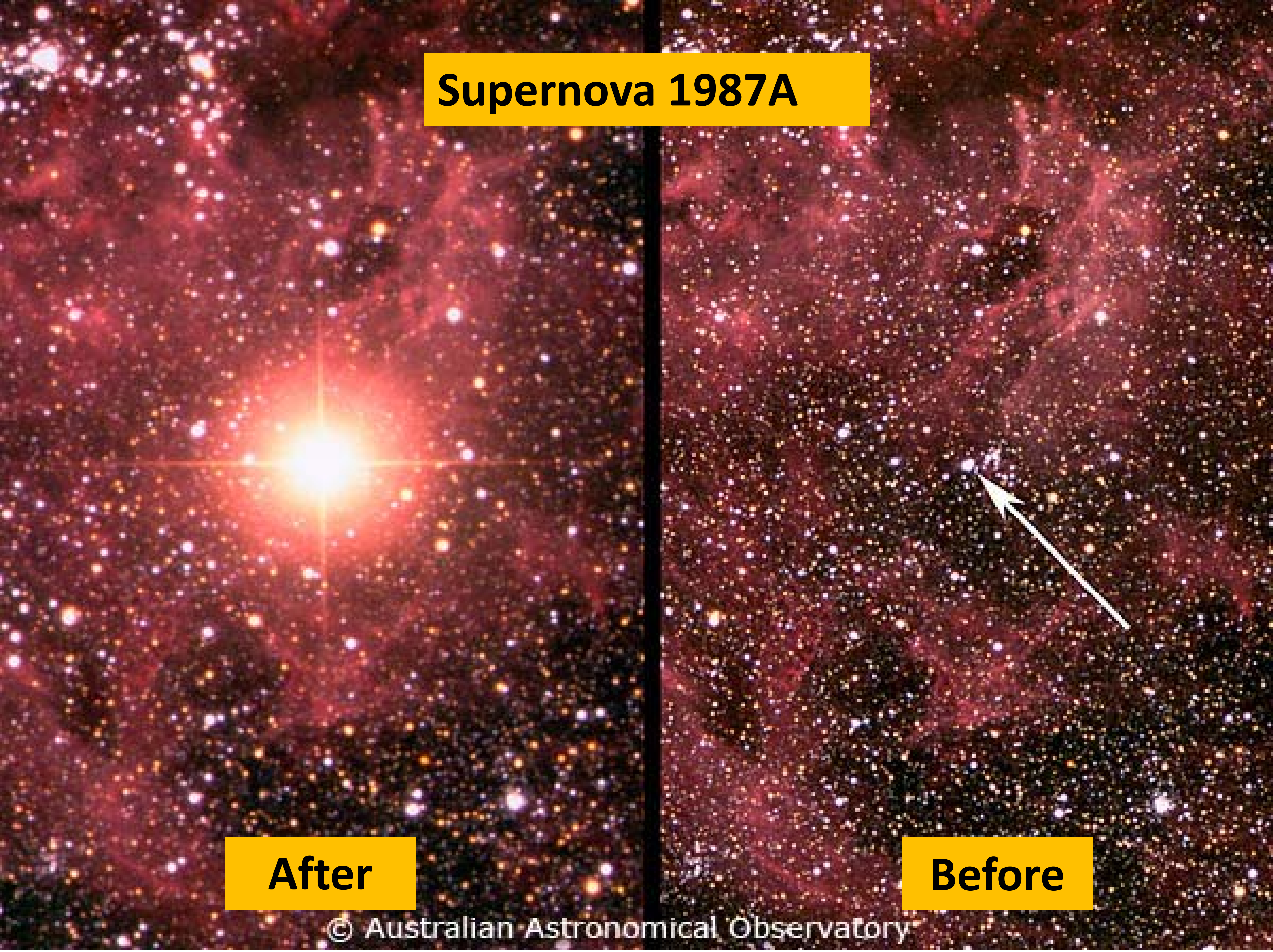
Introduction: *Supernova Explosions*

Gravitational Binding Energy

$$E_g = \frac{3}{5}\frac{G_N M_c^2}{R_{NS}} \approx 3.0\times 10^{53}\,\mathrm{erg}\left(\frac{M_c}{1.5M_{sun}}\right)^2\left(\frac{10\mathrm{km}}{R_{NS}}\right)$$



outshine the host galaxy


Supernova 1987A
After
Before
© Australian Astronomical Observatory

Introduction: *Supernova Explosions*

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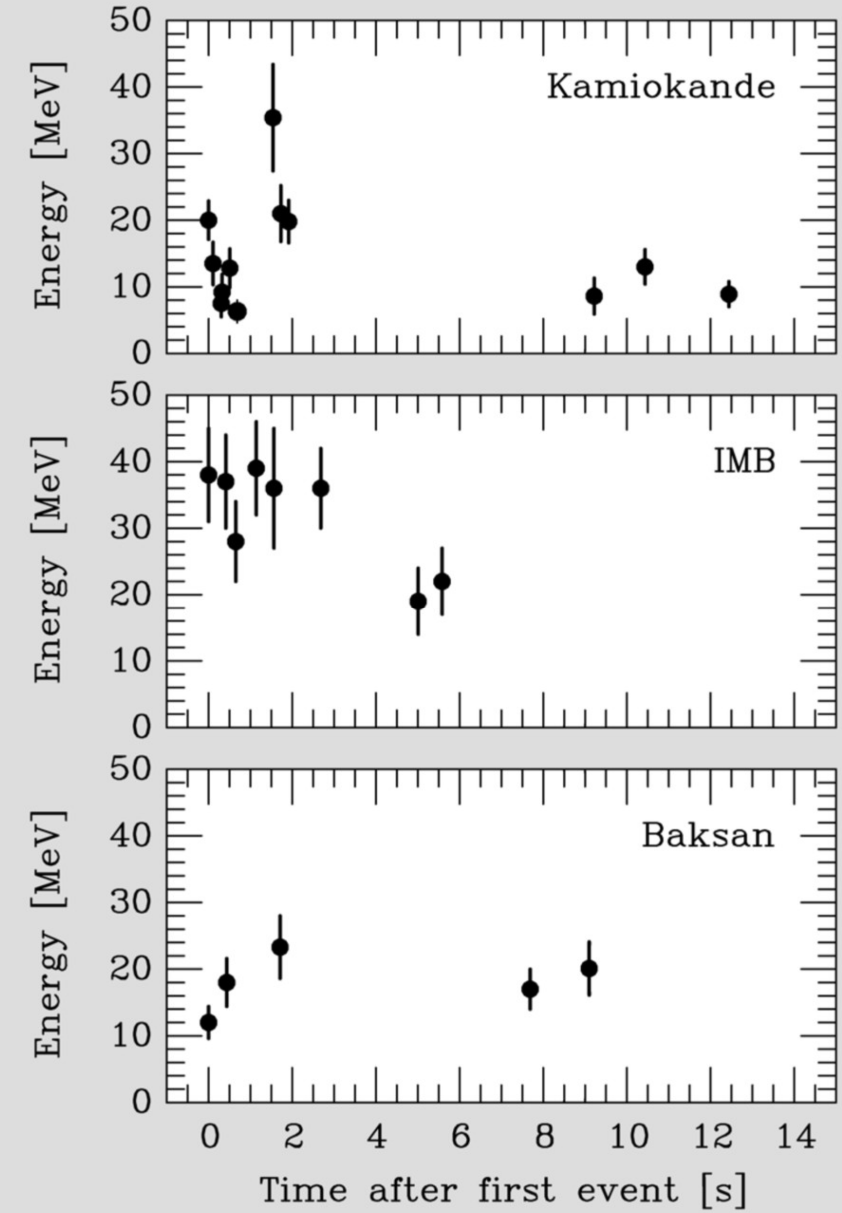
Neutrino Signals from SN 1987A

Kamiokande (Japan) and **IMB** (US):
- Water Cerenkov Detector
- Intended for Proton Decays

Baksan (Soviet Union):
- Scintillator Telescope
- Clock Uncertainty: + 2/-54 s



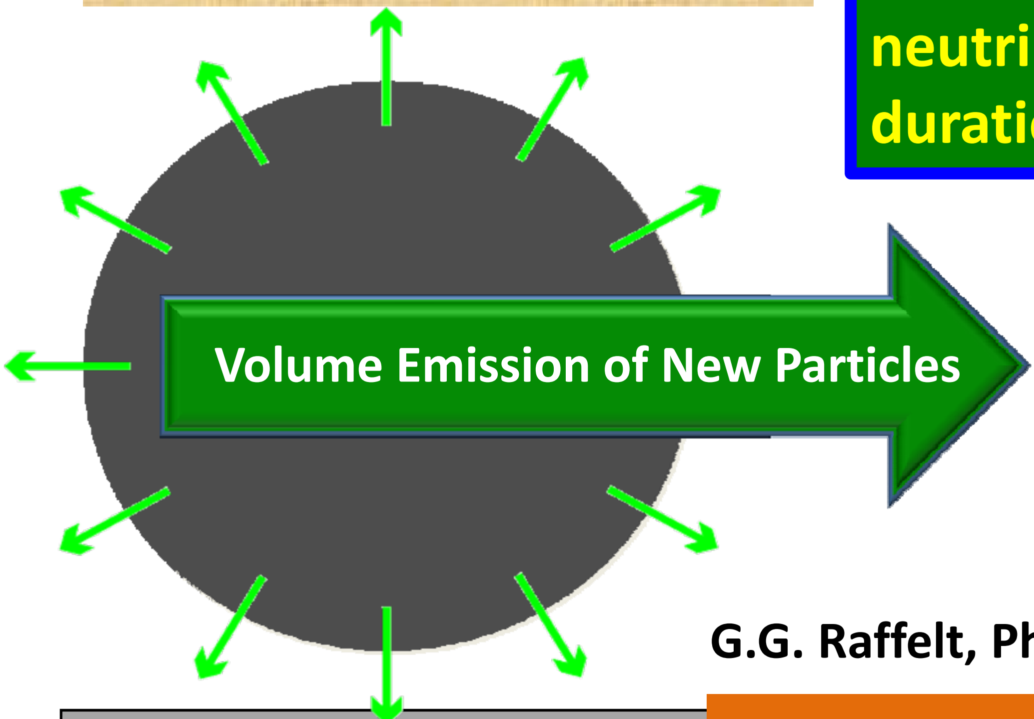
Information from observations

- Standard Picture of Core-collapse SNe
- Neutrino Diffusion Time from the NS
- Neutrino Luminosity $L_\nu = 10^{53}$ erg s^{-1}

Introduction: *Standard energy-loss arguments*

Neutrino Cooling

$$\mathcal{E}_\nu = 3.0 \times 10^{33}\ \mathrm{erg\ cm^{-3}\ s^{-1}}$$



$$\mathcal{E}_{\mathrm{new}} < 3.0 \times 10^{33}\ \mathrm{erg\ cm^{-3}\ s^{-1}}$$

If new weakly interacting particles can be produced in the supernova core, they will steal energies from neutrino bursts, which reduces the duration of neutrino signals.

Proto-Neutron Star

$\rho \approx \rho_{\mathrm{nuc}} = 3 \times 10^{14}\ \mathrm{g\ cm^{-3}}$

$T \approx 30\ \mathrm{MeV}$

G.G. Raffelt, Phys. Rept. 198 (1990) 1

Axions, Majorons, *Sterile Neutrinos, …*

Sterile Neutrinos in SN Cores

Production of sterile neutrinos:

Low matter density: neutrino flavor oscillations with matter effects

High matter density: production & absorption via scattering processes

Occupation-number formalism:

Sigl, Raffelt, 93'

$$\rho_{ij} = \langle b_i^+ b_j \rangle$$

$$\bar{\rho}_{ij} = \langle d_j^+ d_i \rangle$$

Diagonal terms: just the usual occupation numbers

Non-diagonal terms: encode the phase information

Equations of motion:

$$\dot{\rho}_{\mathbf{p}} = i[\rho_{\mathbf{p}}, \Omega_{\mathbf{p}}] + \sum_{i=1}^{n} \left[\left(I_i - \frac{1}{2}\{I_i, \rho_{\mathbf{p}}\} \right) \mathcal{P}_{\mathbf{p}}^i - \frac{1}{2}\{I_i, \rho_{\mathbf{p}}\} \mathcal{A}_{\mathbf{p}}^i \right]$$

Neutral-current interaction

$$+ \frac{1}{2} \sum_a \int \frac{d^3\mathbf{p}}{(2\pi)^3} \left[\mathcal{W}_{\mathbf{p}'\mathbf{p}}^a \left(G^a \rho_{\mathbf{p}'} G^a (1-\rho_{\mathbf{p}}) + \text{h.c.} \right) - \mathcal{W}_{\mathbf{p}\mathbf{p}'}^a \left(\rho_{\mathbf{p}} G^a (1-\rho_{\mathbf{p}'}) G^a + \text{h.c.} \right) \right]$$

Sterile Neutrinos in SN Cores

Two-flavor mixing case

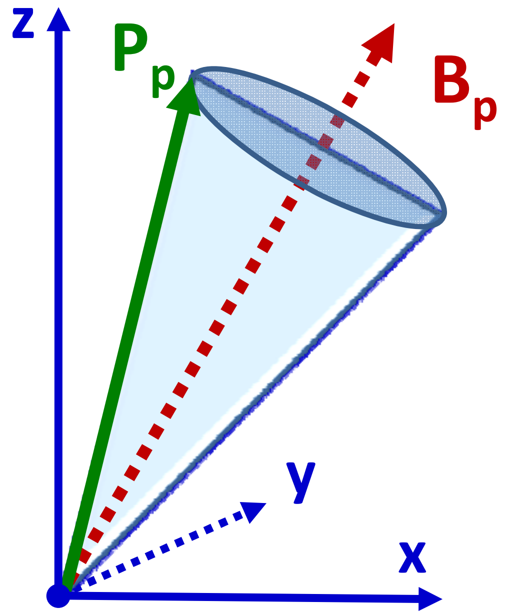
$$\rho_{\mathrm{p}} = \frac{1}{2}\left(n_{\mathrm{p}} + \mathbf{P}_{\mathrm{p}} \cdot \boldsymbol{\tau}\right)$$

$$\Omega_{\mathrm{p}} = \frac{1}{2}\left(E_{\mathrm{p}} + \mathbf{B}_{\mathrm{p}} \cdot \boldsymbol{\tau}\right)$$

$$\dot{\rho}_{\mathrm{p}} = i[\rho_{\mathrm{p}}, \Omega_{\mathrm{p}}] \longrightarrow \dot{\mathbf{P}}_{\mathrm{p}} = \mathbf{B}_{\mathrm{p}} \times \mathbf{P}_{\mathrm{p}}$$

Flavor polarization vectors rotate around magnetic fields

$$\mathbf{B}_{\mathrm{p}} = \left(\frac{\Delta m^2}{2E}\sin 2\vartheta, \quad 0, \quad \frac{\Delta m^2}{2E}\cos 2\vartheta - V_{\mathrm{eff}}\right)$$



Matter Effects

Wolfenstein, 78'; Mikheyev, Smirnov, 85

$$\sin^2 2\vartheta_{\nu,\bar{\nu}} = \frac{\sin^2 2\vartheta}{\sin^2 2\vartheta + \left(\cos 2\vartheta \mp (\pm) E / E_r\right)^2}$$

where the resonant energy is

$$E_r = \frac{\Delta m^2}{2|V_{\mathrm{eff}}|}$$

maximal mixing if $E \sim E_r \cos 2\vartheta$

Sterile Neutrinos in SN Cores

Weak-damping limit

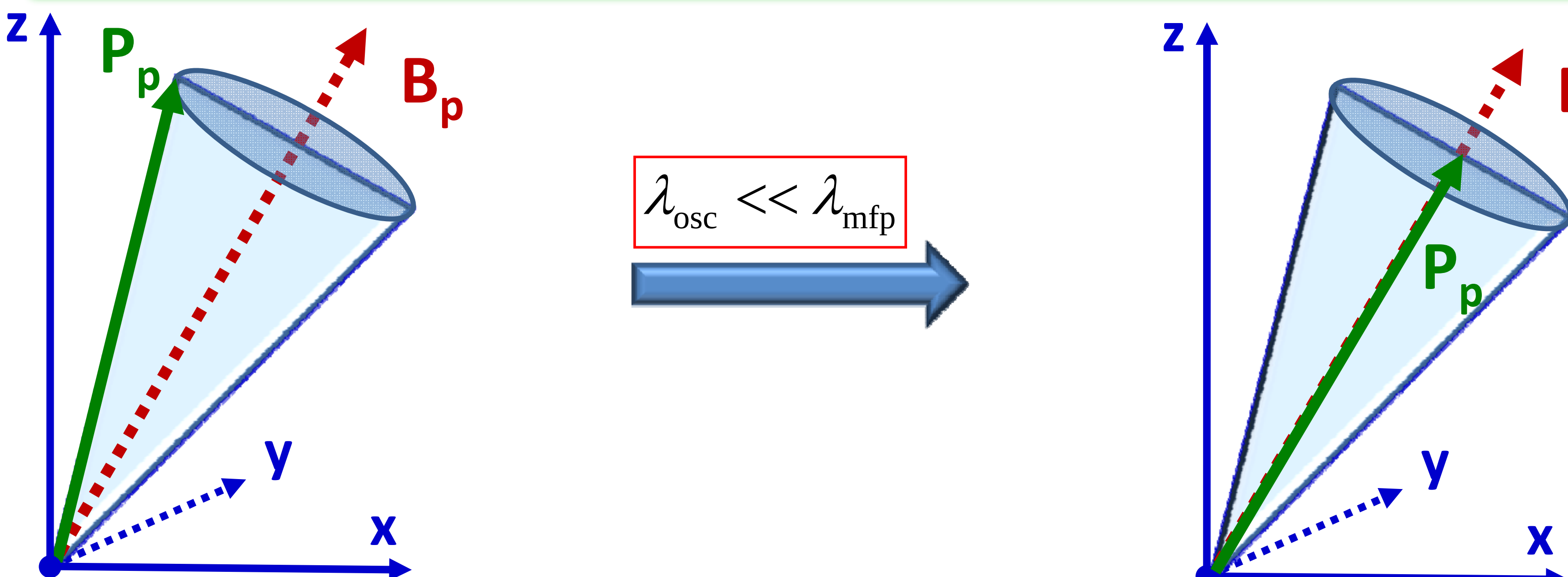
Oscillation length

$$\lambda_{\rm osc} = \frac{4\pi E}{\Delta\tilde{m}^2} < 0.7\,{\rm cm}\left(\frac{E}{30\ {\rm MeV}}\right)\left(\frac{10^{-4}}{\sin 2\vartheta}\right)\left(\frac{10\,{\rm keV}}{m_s}\right)^2$$

Mean free path

$$\lambda_{\rm mfp} = \frac{1}{N_{\rm B}\sigma_{\nu N}} \approx 10^3\,{\rm cm}\left(\frac{30\ {\rm MeV}}{E}\right)^2\left(\frac{10^{14}\,{\rm g\,cm^{-3}}}{\rho}\right)$$

Neutrinos oscillate many times before a subsequent collision with nucleons



Sterile Neutrinos in SN Cores

In the weak-damping limit

$$\boxed{\tilde{\rho}_{\mathbf{p}} = \frac{1}{2}\left[n_{\mathbf{p}} + \left(\mathbf{P}_{\mathbf{p}} \cdot \hat{\mathbf{B}}_{\mathbf{p}}\right)\left(\hat{\mathbf{B}}_{\mathbf{p}} \cdot \boldsymbol{\tau}\right)\right]}$$

averaged over a period of oscillation

Two independent parameters: occupation numbers $f^{\alpha}_{\mathbf{p}}$ & $f^{s}_{\mathbf{p}}$

$$\boxed{\tilde{\rho}_{\mathbf{p}} = \begin{pmatrix} f_{\mathbf{p}}^{\alpha} & 0 \\ 0 & f_{\mathbf{p}}^{s} \end{pmatrix} + \frac{1}{2}\left(f_{\mathbf{p}}^{\alpha} - f_{\mathbf{p}}^{s}\right) t_{\mathbf{p}} \begin{pmatrix} 0 & 1 \\ 1 & 0 \end{pmatrix}}$$

Simplified equations of motion:

$$\dot{f}_{\mathbf{p}}^{s} = \frac{1}{4} s_{\mathbf{p}}^{2} \left\{ \left[(1 - f_{\mathbf{p}}^{s}) \mathscr{P}_{\mathbf{p}}^{\alpha} - f_{\mathbf{p}}^{s} \mathscr{A}_{\mathbf{p}}^{\alpha}\right] + \sum_{a} \left(g_{\alpha}^{a}\right)^{2} \int \frac{d^{3}\mathbf{p}}{(2\pi)^{3}} \left[\mathscr{W}_{\mathbf{p}'\mathbf{p}}^{a} f_{\mathbf{p}'}^{\alpha} (1 - f_{\mathbf{p}}^{s}) - \mathscr{W}_{\mathbf{p}\mathbf{p}'}^{a} f_{\mathbf{p}}^{s} (1 - f_{\mathbf{p}'}^{\alpha})\right] \right\}$$

further simplification if sterile neutrinos escape from the SN core

$$\dot{f}_{\mathbf{p}}^{s} = \frac{1}{4} s_{\mathbf{p}}^{2} \left[\mathscr{P}_{\mathbf{p}}^{\alpha} + \sum_{a} \left(g_{\alpha}^{a}\right)^{2} \int \frac{d^{3}\mathbf{p}}{(2\pi)^{3}} \mathscr{W}_{\mathbf{p}'\mathbf{p}}^{a} f_{\mathbf{p}'}^{\alpha} \right]$$

set $f^{s}_{\mathbf{p}} = 0$

Lepton-number-loss rate $\dot{\mathscr{N}}_{\mathrm{L}} = \int \frac{d^{3}\mathbf{p}}{(2\pi)^{3}} \dot{f}_{\mathbf{p}}^{s}$

Energy-loss rate $\mathcal{E}_{s} = \int \frac{d^{3}\mathbf{p}}{(2\pi)^{3}} E \dot{f}_{\mathbf{p}}^{s}$

Sterile Neutrinos in SN Cores

Neutrino matter potentials

$$V_{\nu_e} = \sqrt{2} G_{\mathrm{F}} N_{\mathrm{B}} \left[Y_e - \frac{1}{2} Y_n + 2Y_{\nu_e} + Y_{\nu_\mu} + Y_{\nu_\tau} \right]$$

$$V_{\nu_\mu} = \sqrt{2} G_{\mathrm{F}} N_{\mathrm{B}} \left[-\frac{1}{2} Y_n + Y_{\nu_e} + 2Y_{\nu_\mu} + Y_{\nu_\tau} \right]$$

$$\boxed{V_{\nu_\tau} = \sqrt{2} G_{\mathrm{F}} N_{\mathrm{B}} \left[-\frac{1}{2} Y_n + Y_{\nu_e} + Y_{\nu_\mu} + 2Y_{\nu_\tau} \right]}$$

Remarks:

1. degenerate electron neutrinos; the equation of state involved; charged current interactions; so we consider tau neutrinos for simplicity;
2. we assume the SN core to be homogeneous and isotropic.

Tau-sterile neutrino mixing

$$V_{\nu_\tau} = -\frac{G_{\mathrm{F}}}{\sqrt{2}} N_{\mathrm{B}} \left(1 - Y_e - 2Y_{\nu_e} - 4Y_{\nu_\tau}\right) < 0$$

Initial conditions:

$Y_e = 0.3$, $Y_{\nu_e} = 0.07$, $Y_{\nu_\mu} = Y_{\nu_\tau} = 0$

$$\sin^2 2\vartheta_{\nu,\bar{\nu}} = \frac{\sin^2 2\vartheta}{\sin^2 2\vartheta + \left(\cos 2\vartheta \pm E / E_r\right)^2}$$

1. the MSW resonance occurs in the antineutrino channel;
2. asymmetry between tau neutrinos and antineutrinos.

Energy-loss Rates and SN Bounds

Simple bounds in the 'vacuum limit': $E_r >> E$

$$\sin^2 2\vartheta_{\nu,\bar{\nu}} = \frac{\sin^2 2\vartheta}{\sin^2 2\vartheta + \left(\cos 2\vartheta \pm E/E_r\right)^2} \Longrightarrow \boxed{\vartheta_\nu \approx \vartheta_{\bar{\nu}} \approx \vartheta}$$

Energy-loss rates

$$\mathcal{E}_s = \underbrace{\boxed{2}}_{\nu+\bar{\nu}} \int_0^\infty \frac{E^2}{2\pi^2} \frac{E}{\exp(E/T)+1} \left(\frac{1}{4}\sin^2 2\vartheta\right) \underbrace{\boxed{\frac{N_B G_F^2 E^2}{\pi}}}_{\nu - N} dE = 4 N_B G_F^2 T^6 \vartheta^2$$

Supernova Bound

$$\mathcal{E}_s = 4 N_B G_F^2 T^6 \vartheta^2 < \mathcal{E}_\nu = 3.0 \times 10^{33}\ \mathrm{erg\,cm^{-3}s^{-1}} \Longrightarrow \boxed{\vartheta^2 \le 10^{-8}}$$

$\rho \approx \rho_{nuc} = 3 \times 10^{14}\ \mathrm{g\ cm^{-3}}$

$T \approx 30$ MeV

Such a simple bound is valid and mass-independent only in the 'vacuum limit'.

Energy-loss Rates and SN Bounds

Conditions for the vacuum limit

$$E_r = 3.25\,\mathrm{MeV}\left(\frac{m_s}{10\,\mathrm{keV}}\right)^2\left(\frac{10^{14}\,\mathrm{g\,cm^{-3}}}{\rho}\right)\left|Y_0 - Y_{\nu_\tau}\right|^{-1}$$

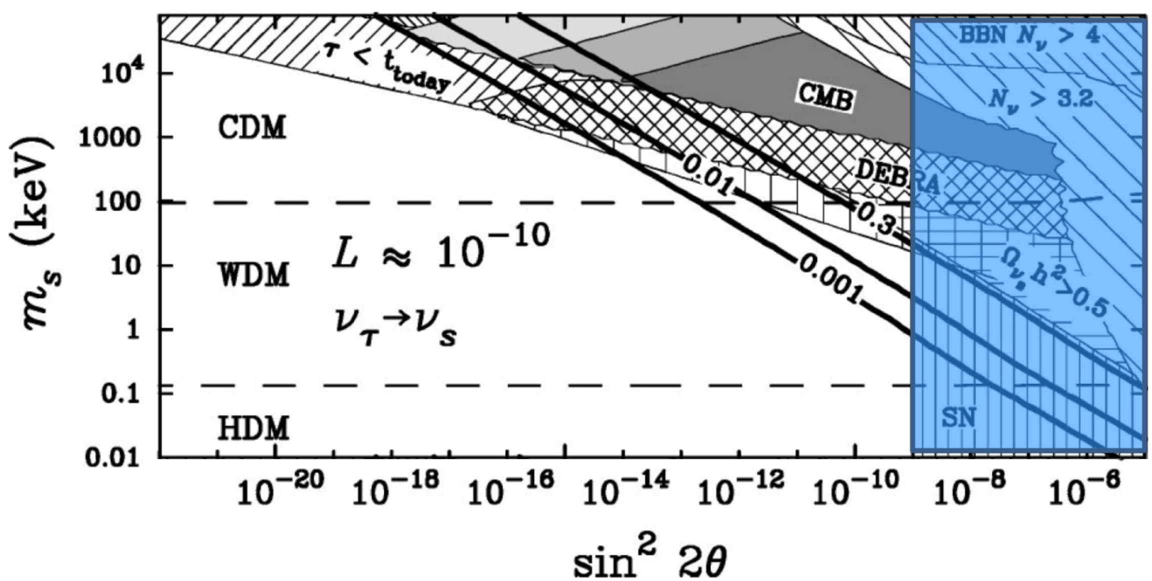
$Y_0 = (1 - Y_e - 2Y_{\nu_e})/4 = 0.14$

$\boxed{m_s > 100\,\mathrm{keV}}$

$\boxed{Y_{\nu_\tau} \to Y_0}$

Is the simple bound also valid for the small-mass range?

Abazajian, Fuller, Patel, 01'



A Stationary State?

$\boxed{Y_{\nu_\tau} \to Y_0}$ ⇒ $\boxed{\vartheta_\nu \approx \vartheta_{\bar\nu} \approx \vartheta}$

Including the degen. param.

$$f_E^{\nu_\tau} = \frac{1}{\exp[E/T - \eta] + 1}$$

$$f_E^{\bar\nu_\tau} = \frac{1}{\exp[E/T + \eta] + 1}$$

Energy-loss Rates and SN Bounds

Criterion for a stationary state

Emission Rate = Mixing angle in matter × Population of active neutrinos × Neutral-current interactions

Neutrino Emission Rate = Antineutrino Emission Rate

Evolution of the degeneracy parameter

$$\dot{N}_{\nu_\tau} = -\frac{1}{4}\sum_a \int \frac{E^2 dE}{2\pi^2} \sin^2 2\vartheta_\nu \int \frac{E'^2 dE'}{2\pi^2} w^a_{E'E} f^{\nu_\tau}_{E'}$$

$$\dot{N}_{\overline{\nu}_\tau} = -\frac{1}{4}\sum_a \int \frac{E^2 dE}{2\pi^2} \sin^2 2\vartheta_{\overline{\nu}} \int \frac{E'^2 dE'}{2\pi^2} \overline{w}^a_{E'E} f^{\overline{\nu}_\tau}_{E'}$$

$$f^{\nu_\tau}_E = \frac{1}{\exp[E/T-\eta]+1}$$

$$f^{\overline{\nu}_\tau}_E = \frac{1}{\exp[E/T+\eta]+1}$$

Sterile neutrinos with mixing angles $\vartheta_\nu < \vartheta_c \approx 10^{-2}$ can escape from the core.

Energy-loss Rates and SN Bounds

Evolution of the degeneracy parameter

$$\frac{d}{dt}\eta(t) = \frac{N_{\mathrm{B}} G_{\mathrm{F}}^2 s_{2\vartheta}^2 T^2}{4\pi}\left[\mathcal{F}_{\bar{\nu}}(\eta) - \mathcal{F}_{\nu}(\eta)\right]\boxed{\mathcal{G}^{-1}(\eta)}$$

$$\mathcal{G}(\eta) = \frac{d}{d\eta}\left[F_2(\eta) - F_2(-\eta)\right]$$

Fermi-Dirac Integral $$F_2(\eta) = \int_0^\infty \frac{x^2}{e^{x-\eta}+1}dx$$

Energy-loss Rates and SN Bounds

Evolution of the degeneracy parameter

$$\frac{d}{dt}\eta(t) = \frac{N_B G_F^2 s_{2\vartheta}^2 T^2}{4\pi}\left[\mathcal{F}_{\bar{\nu}}(\eta) - \mathcal{F}_{\nu}(\eta)\right]\mathcal{G}^{-1}(\eta)$$

Feedback effects

Initial condition: *t = 0, η = 0*

$\boxed{\mathcal{F}_{\bar{\nu}}(0) - \mathcal{F}_{\nu}(0) > 0}$ **Antineutrino Emission Rate > Neutrino Emission Rate**

η increases (η >0)

Y_{ν_τ} **(τ number asymmetry)** ↑

$E_r \propto \left|Y_0 - Y_{\nu_\tau}\right|^{-1}$ ↑

$\vartheta_{\bar{\nu}}$ ↓ ϑ_{ν} ↑

$$f_E^{\nu_\tau} = \frac{1}{\exp[E/T - \eta] + 1}$$ ↑

$$f_E^{\bar{\nu}_\tau} = \frac{1}{\exp[E/T + \eta] + 1}$$ ↓

Antineutrino Emission Rate ↓ **Neutrino Emission Rate** ↑

stable point η* $\boxed{\mathcal{F}_{\bar{\nu}}(\eta^*) = \mathcal{F}_{\nu}(\eta^*)}$

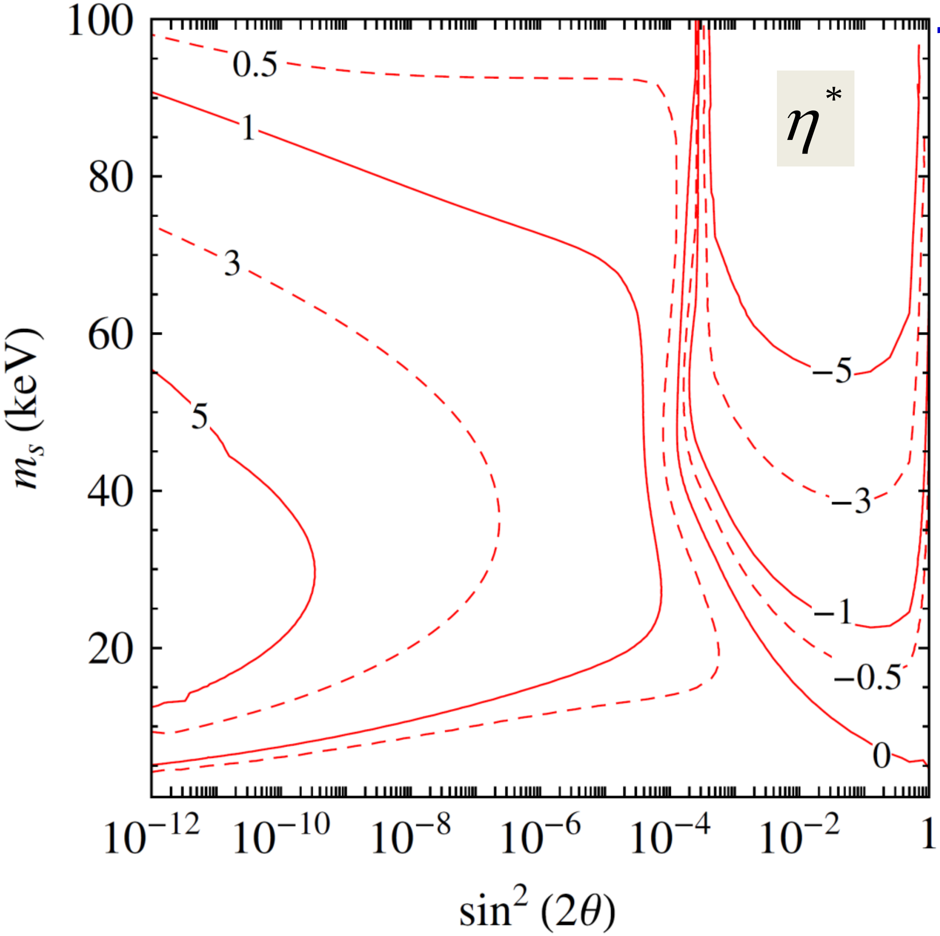
Neutrino Emission Rate = Antineutrino Emission Rate

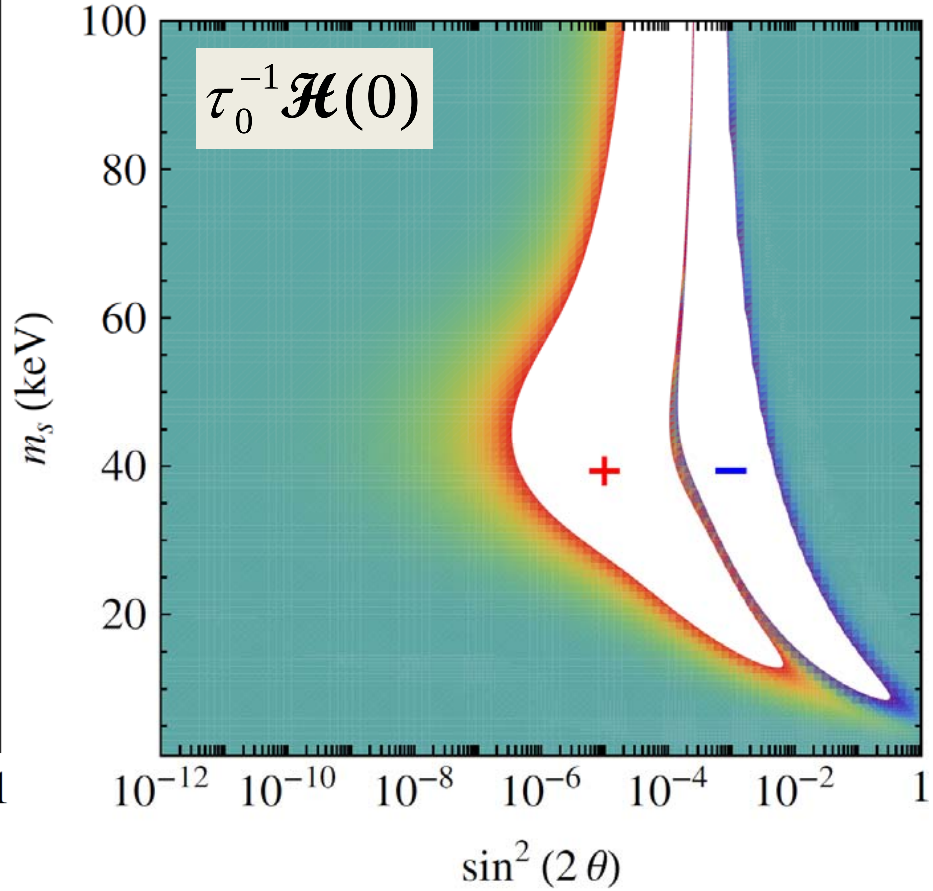
Energy-loss Rates and SN Bounds

Evolution of the degeneracy parameter

Raffelt, Zhou, 11'

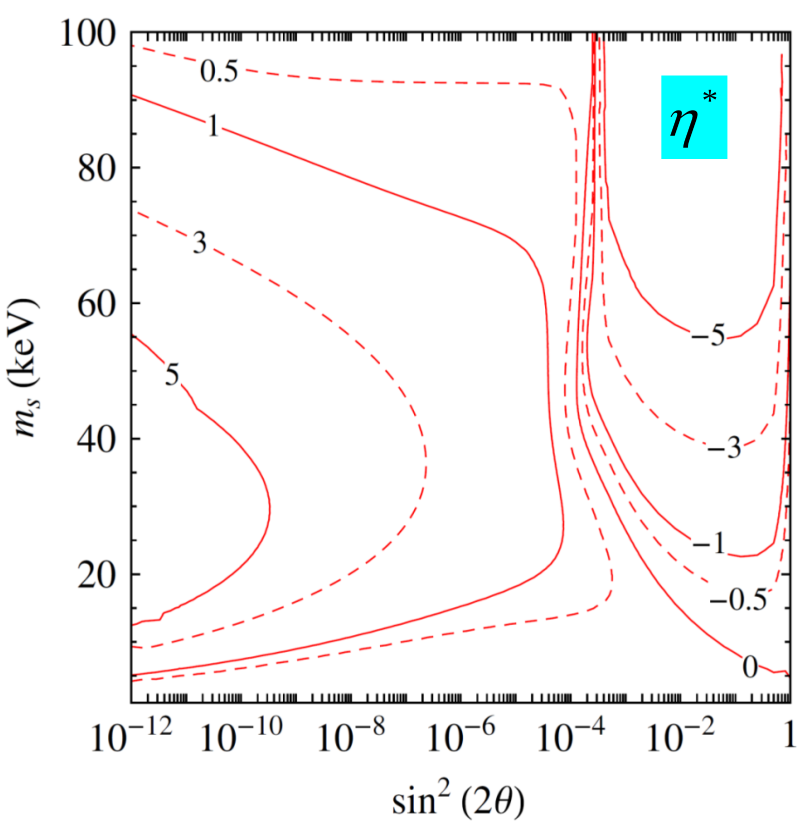
$$\frac{d}{dt}\eta(t) = \boxed{\frac{N_{\rm B} G_{\rm F}^2 s_{2\vartheta}^2 T^2}{4\pi}} \boxed{\left[\mathcal{F}_{\bar\nu}(\eta) - \mathcal{F}_{\nu}(\eta)\right]\mathcal{G}^{-1}(\eta)}$$

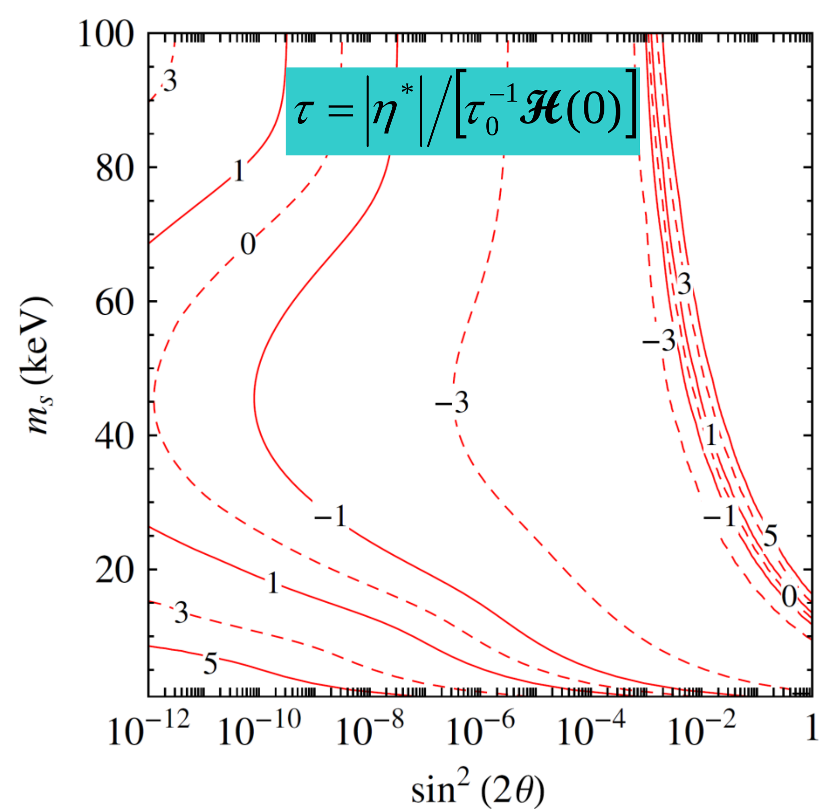




Energy-loss Rates and SN Bounds

1. The stable point η^* can be either negative or positive, depending on the sterile neutrino mass and vacuum mixing angle;
2. The values of η^* are negative for large vacuum mixing angles, because more antineutrinos than neutrinos are trapped in the SN core;
3. We temporarily ignore the trapped sterile neutrinos, which may actually transfer energies rapidly due to their larger mean free paths.





Energy-loss Rates and SN Bounds

Raffelt, Zhou, 11'

Energy-loss rate

$$\mathcal{E}_s(t) = \frac{N_B G_F^2 s_{2\vartheta}^2 T^6}{8\pi^3}\left[\mathcal{R}_{\bar{\nu}}(\eta) + \mathcal{R}_{\nu}(\eta)\right]$$

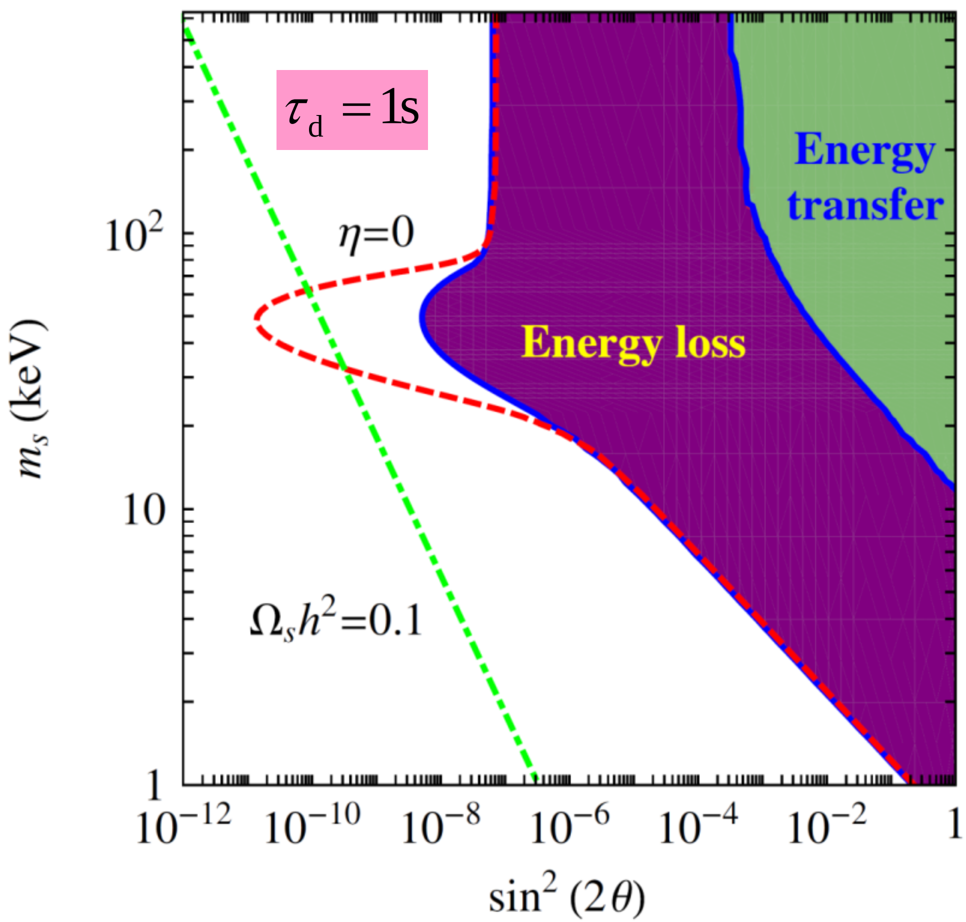
$$\mathcal{R}_{\bar{\nu}}(\eta) = \int_0^\infty \frac{x^5}{e^{x+\eta}+1}\frac{1-\mathcal{B}(x, x_r\varepsilon^-, x_r\varepsilon^+)}{s_{2\vartheta}^2 + (c_{2\vartheta} - x/x_r)^2}dx$$

Averaged energy-loss rate

$$\boxed{\langle \mathcal{E}_s \rangle = \tau_d^{-1}\int_0^{\tau_d} \mathcal{E}_s(t)dt}$$

Supernova bound

$$\boxed{\langle \mathcal{E}_s \rangle < 3.0\times 10^{33}\,\mathrm{erg\,cm^{-3}s^{-1}}}$$



Sterile Neutrinos and SN Explosions

Mixing with electron neutrinos

$$e^- + p \Leftrightarrow n + \nu_e$$

ν_s escape from the SN cores

$$\boxed{\mu_e} - \boxed{\mu_{\nu_e}} = \boxed{\mu_n - \mu_p} + \delta m_{np}$$

Equation of state

How to constrain sterile neutrinos?

Remarks:

1. If the lepton-number loss is not significant, one can simply apply the standard energy-loss argument to the ν_e-ν_s mixing case;
2. For the warm-dark-matter mass range (*1* keV to *10* keV), the MSW resonance may be present and amplify the lepton-number-loss rate;
3. Sterile neutrinos have already done something important during the collapsing phase, such as reducing the electron number fraction Y_e and thus the size of the homologous core, and the energy of the shock wave.

Sterile Neutrinos and SN Explosions

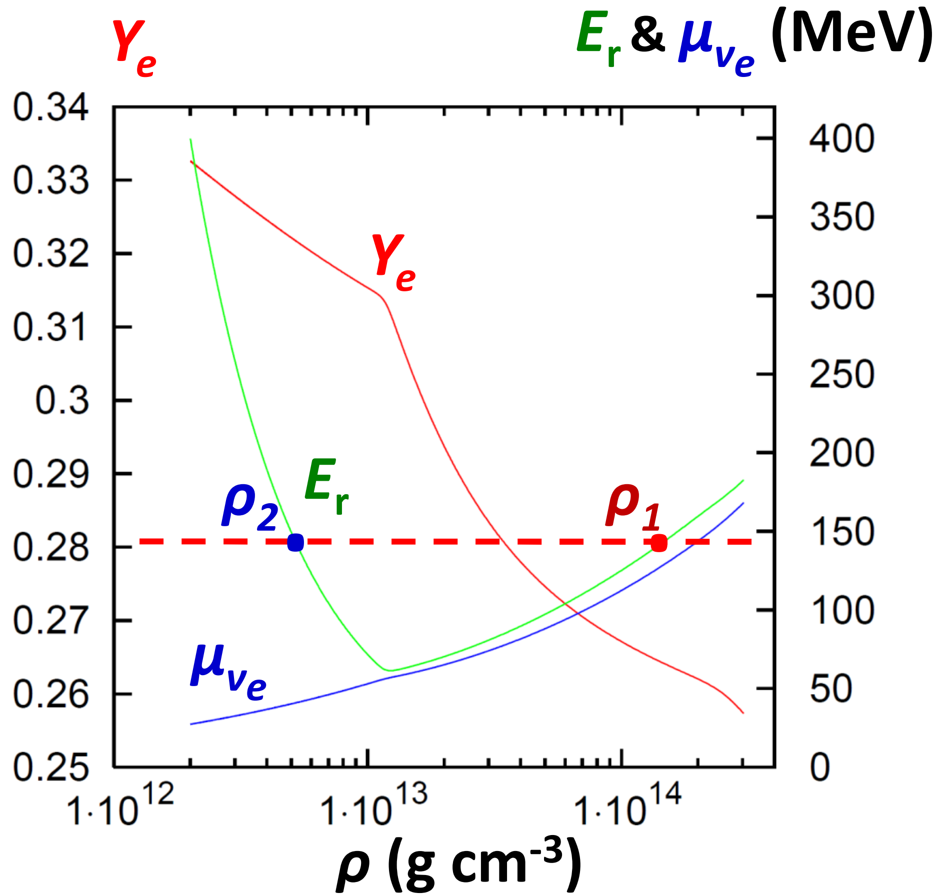
Sterile neutrino assisted SN explosions?

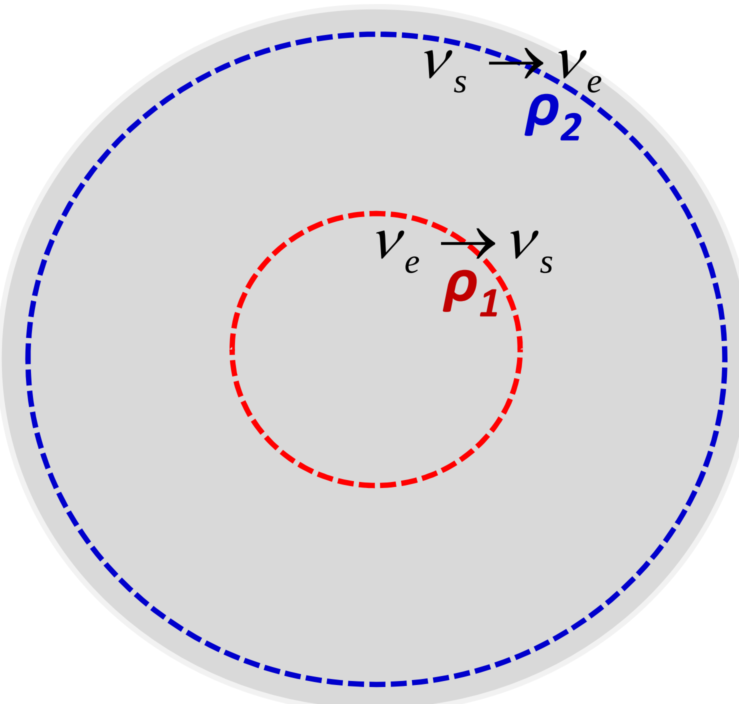
Hidaka, Fuller, 06’

One-zone model of the collapsing core: the EoS & resonant ν_e-ν_s conversion,…

To include the neutrino trapping and diffusion, shock-wave propagation, …

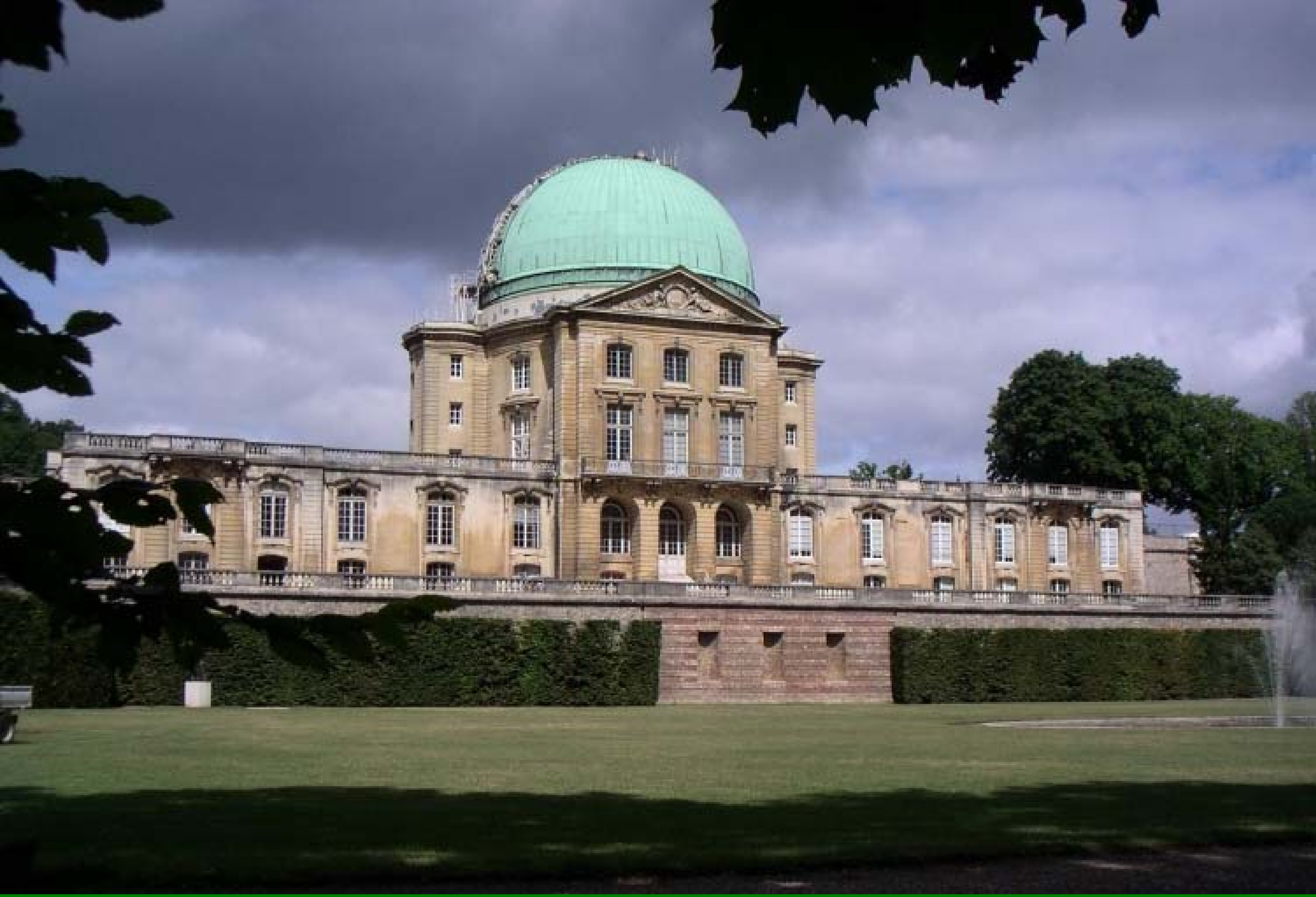
Raffelt, Zhou, work in progress





Summary

1. **Sterile neutrinos of keV masses are a promising candidate for warm dark matter. They are expected to fill up the 'desert' in the fermion mass spectrum.**



Thank you for your attention!