

## Quantum warm dark matter fermions and gravitation determine the observed galaxy structures

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Dark matter (DM) is the main component of galaxies. Quantum mechanics is a cornerstone of physics from microscopic to macroscopic systems as quantum liquids  $\text{He}^3$ , white dwarf stars and neutron stars. A study by a team of CNRS scientists from Paris Observatory, Pierre et Marie Curie University, and INFN-Milano Bicocca, suggests for the first time that quantum mechanics is also responsible of galaxy structures at the kpc scales and below: near the galaxy center, below  $10 - 100$  pc, the *quantum* effects are important for warm dark matter (WDM), that is for dark matter particles with masses in the keV scale. A new approach to galaxy structure and a set of results follow in remarkable agreement with observations: (i) Dwarf galaxies turn to be quantum macroscopic objects for WDM supported against gravity by the WDM fermion pressure. (ii) A theoretical *analytic* framework from atomic physics is implemented to determine galaxy structure from the most compact dwarf galaxies to the largest dilute galaxies (spirals, ellipticals). The obtained galaxy mass, halo radius, phase-space density, velocity dispersion, are fully consistent with observations. (iii) Interestingly enough, a *minimal* mass and *minimal* velocity dispersion are found for DM dominated objects, which in turn imply an universal minimal mass  $m_{min} = 1.9$  keV for the WDM particle.

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Dark matter (DM) is the main component of galaxies, especially of dwarf galaxies. Cold DM (CDM) is formed by very heavy particles, which are very slow (almost zero velocity), and produce a huge overabundance of small structures, substructures or "satellites" below  $\sim 50$  kpc till very small scales  $\sim 0.005$  pc. This constitutes, as is well known, one of the most serious drawbacks for CDM (also known as the "CDM satellites problem"). On the contrary, Warm DM (WDM) is formed by particles with masses in the intermediate scales, the keV scale, which have high enough velocities and produce DM structures in the range of scales  $\lesssim 50$  kpc in agreement with observations: In WDM structure formation, substructures below the scale  $\sim 50$  kpc are naturally not formed. For all scales larger than 50 kpc, WDM yields the same results than CDM and agrees with the galactic as well as cosmological observations: small as well as large scale structure observations and CMB anisotropies.

Astronomical observations show that the DM galaxy density profiles are cored, that is, profiles which are flat at the center. On the contrary,  $N$ -body CDM simulations exhibit cusped density profiles, with a typical  $1/r$  cusped behaviour near the galaxy center  $r = 0$ . Classical  $N$ -body WDM simulations exhibit cores but with sizes much smaller than the observed cores. A new approach to this problem has been recently developed by the team thanks to quantum mechanics.

Fermions always provide a non vanishing pressure of quantum nature due to the combined action of the Pauli exclusion principle and Heisenberg uncertainty principle. Quantum effects for WDM fermions rule out the presence of galaxy cusps for WDM and enlarge the classical core sizes because their *repulsive and non-local* nature extend well beyond the small pc scales. Smoothing the density profile at the central regions has an effect on the whole galaxy halo.

The team developed a physical *analytic* framework to the galaxy structures which accounts for the quantum nature of the WDM fermions and determines selfconsistently the gravitational potential and the *galaxy equation of state*. The *full range* of physical galaxy situations is covered: from the compact dwarfs (in the extreme quantum degenerate fermions limit) till the large galaxies, spiral and ellipticals (in the dilute classical and semiclassical regime). In this framework, some results obtained by the team are the following:

- A one parameter family of galaxy solutions parametrized by the value of the chemical potential at the origin  $\nu(0)$  or equivalently, by the phase-space density at the origin  $Q(0)$ : Large positive values of  $\nu(0)$ , or  $Q(0)$ , correspond to the most compact galaxies while large negative values of  $\nu(0)$  yield dilute objects (intermediate and large galaxies, spiral and ellipticals). Approaching the dilute regime yields larger and larger halo radii, galaxy masses and velocity dispersions. Their maximum values are limited by the initial conditions provided by the primordial power spectrum.
- All obtained density profiles are cored. The core sizes  $r_h$  are in agreement with the observations, from the compact galaxies where  $r_h \sim 20$  pc till the spiral and elliptical galaxies where  $r_h \sim 0.2 - 60$  kpc. The larger and positive is the chemical potential  $\nu(0)$ , the smaller is the core. The minimal one arises in the degenerate case  $\nu(0) \rightarrow +\infty$  (compact dwarf galaxies). Fig. 1 displays the theoretically obtained density and velocity profiles.

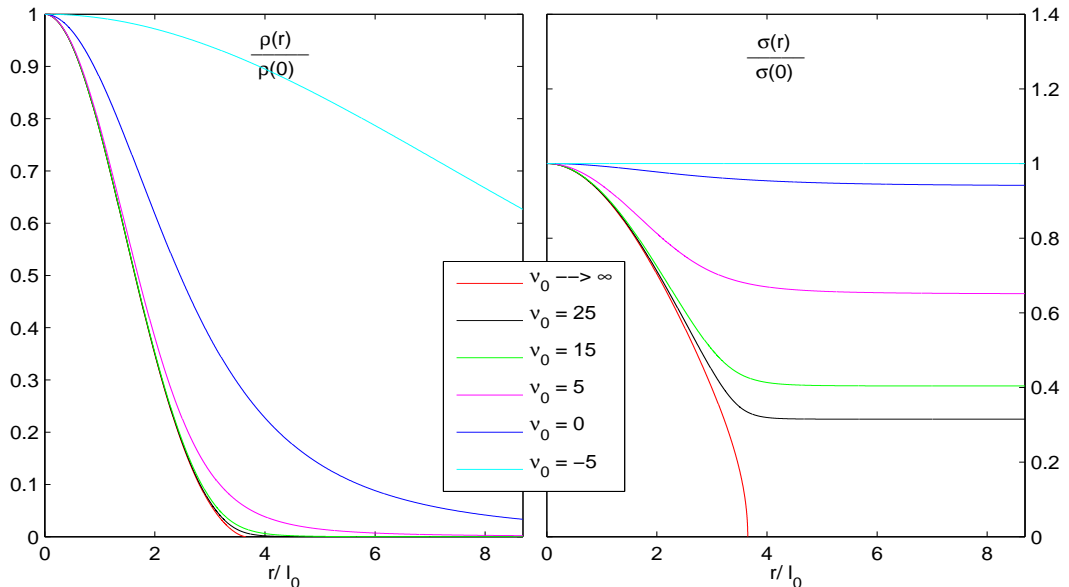


FIG. 1: The obtained density and velocity profiles,  $\rho(r)/\rho_0$  and  $\sigma(r)/\sigma(0)$ , for different values of the chemical potential at the origin  $\nu_0$ . Large positive values of  $\nu_0$  correspond to compact galaxies, while negative values of  $\nu_0$  correspond to the classical regime describing spiral and elliptical galaxies. All density profiles are cored. The sizes of the cores  $r_h$  are in agreement with the observations, from the compact galaxies where  $r_h \sim 35$  pc till the spiral and elliptical galaxies where  $r_h \sim .2 - 60$  kpc. The larger and positive is  $\nu_0$ , the smaller is the core. The minimal one arises in the degenerate case  $\nu_0 \rightarrow +\infty$  (compact dwarf galaxies).

The left panel of fig. 2 shows the obtained (dimensionless) phase-space density  $\hbar^3 Q(0)/(\text{keV})^4$ . The observed values are also depicted. The right panel shows the obtained galaxy halo masses. Good agreement is found in **all** the range of galaxies for a DM particle mass  $m$  around 2 keV. Error bars of the observational data are not shown but they are at least about 10 – 20%.

- A *minimal* galaxy mass and *minimal* velocity dispersion are found. This in turn implies a minimal mass  $m_{min}$  for the WDM particle. This minimal WDM mass is a *universal* value, *independent* of the WDM particle physics model because only relies on the degenerate quantum fermion state, which is universal whatever is the non-degenerate regime. These results and the observed halo radius and mass of the compact galaxies also provide further indication that the WDM particle mass  $m$  is approximately around 2 keV. More precise data will make this estimation more precise.

To **conclude**, it is highly remarkable that in the context of warm dark matter, the quantum description provided by this semiclassical framework, (quantum WDM and classical gravitation), is able to reproduce such broad variety of galaxies, as shown in Figs 1 and 2. The resulting galaxy halo radius, galaxy masses and velocity dispersion are fully consistent with observations for all different types of galaxies, Figs 1, 2. Fermionic WDM treated quantum mechanically is able to reproduce the observed galactic cores and their sizes. In addition, WDM simulations produce the right DM structures in agreement with observations for scales  $\gtrsim$  kpc.

Baryons have not yet included in the present study. This is fully justified for dwarf compact galaxies which are composed today 99.99% of DM. In large galaxies the baryon fraction can reach values up to 1 - 3 %. Fermionic WDM by itself produces galaxies and structures in agreement with observations for all types of galaxies, masses and sizes. Therefore, the effect of including baryons is expected to be a small correction to these pure WDM results, consistent with the fact that dark matter is in average six times more abundant than baryons.

## Reference

Fermionic warm dark matter produces galaxy cores in the observed scales,  
C. Destri, H. J. de Vega, N. G. Sanchez, arXiv:1204.3090,  
New Astronomy, in press.

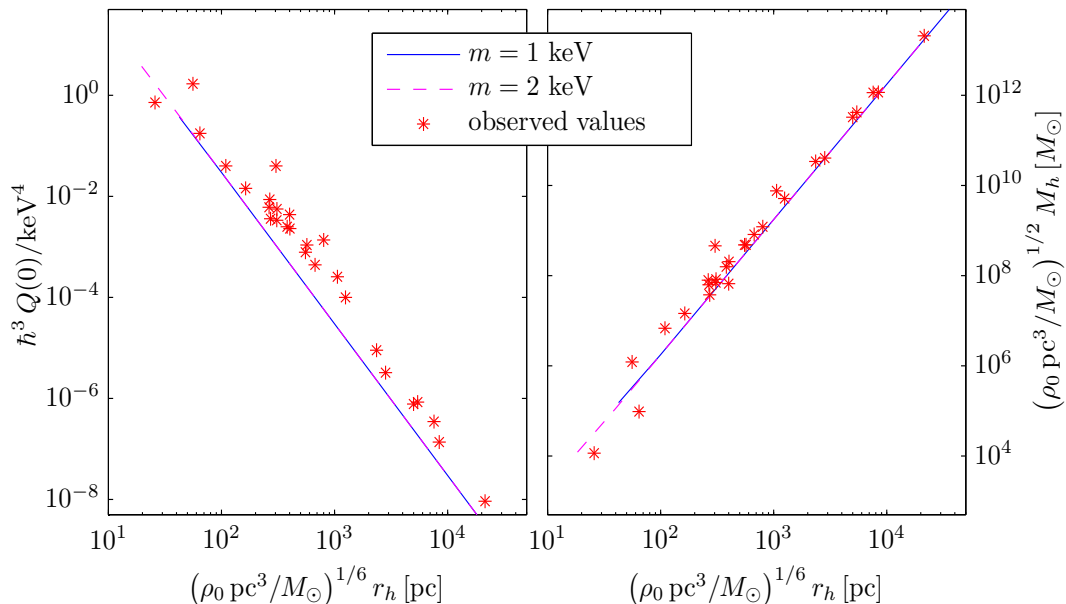


FIG. 2: The left panel displays the (dimensionless) obtained galaxy phase-space density  $\hbar^3 Q(0)/(\text{keV})^4 = (m/\text{keV})^4$  for WDM fermions of mass  $m = 1$  and  $2$  keV versus the ordinary logarithm of the product  $\log_{10}\{r_h [\text{pc}^3 \rho_0/M_\odot]^{1/6}\}$  in parsecs. The red stars  $*$  are the observed values of  $\hbar^3 Q(0)/(\text{keV})^4$ . The observed values  $Q_h$  from the stars' velocity dispersion are in fact upper bounds for the DM  $Q_h$  and therefore the theoretical curve is slightly below them. The right panel displays the obtained galaxy masses  $(M/M_\odot)\sqrt{M_\odot/[\rho_0 \text{ pc}^3]} = 0.82296 \cdot 10^5 (m/\text{keV})^4$  for WDM fermions of mass  $m = 1$  and  $2$  keV versus the product  $r_h [\text{pc}^3 \rho_0/M_\odot]^{1/6}$  in parsecs. The red stars  $*$  are the observed values of  $(M/M_\odot)\sqrt{M_\odot/[\rho_0 \text{ pc}^3]}$ .

Available online :

<http://fr.arxiv.org/abs/1204.3090>

<http://www.sciencedirect.com/science/article/pii/S1384107612001200>

#### Further reading:

The mass of the dark matter particle from theory and observations,  
H. J. de Vega, P. Salucci, N. G. Sanchez,  
New Astronomy **17**, 653 (2012).

Cosmological evolution of warm dark matter fluctuations I: Efficient computational framework with Volterra integral equations,  
H. J. de Vega, N. G. Sanchez, Phys. Rev. **D85**, 043516 (2012)

Cosmological evolution of warm dark matter fluctuations II: Solution from small to large scales and keV sterile neutrinos,  
H. J. de Vega, N. G. Sanchez, Phys. Rev. **D85**, 043517 (2012)

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